

MOSFIRE ABSORPTION LINE SPECTROSCOPY OF $Z > 2$ QUIESCENT GALAXIES: PROBING A PERIOD OF RAPID SIZE GROWTH

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ABSTRACT

Using the MOSFIRE near-infrared multi-slit spectrograph on the Keck 1 Telescope, we have secured high signal-to-noise ratio absorption line spectra for six massive galaxies with redshift $2 < z < 2.5$. Five of these galaxies lie on the red sequence and show signatures of passive stellar populations in their rest-frame optical spectra. By fitting broadened spectral templates we have determined stellar velocity dispersions and, with broad-band HST and Spitzer photometry and imaging, stellar masses and effective radii. Using this enlarged sample of galaxies we confirm earlier suggestions that quiescent galaxies at $z > 2$ have small sizes and large velocity dispersions compared to local galaxies of similar stellar mass. The dynamical masses are in very good agreement with stellar masses ($\log M_*/M_{\text{dyn}} = -0.02 \pm 0.03$), although the average stellar-to-dynamical mass ratio is larger than that found at lower redshift (-0.23 ± 0.05). By assuming evolution at fixed velocity dispersion, not only do we confirm a surprisingly rapid rate of size growth but we also consider the necessary evolutionary track on the mass-size plane and find a slope $\alpha = d \log R_e / d \log M_* \gtrsim 2$ inconsistent with most numerical simulations of minor mergers. Both results suggest an additional mechanism may be required to explain the size growth of early galaxies.

Subject headings: galaxies: evolution — galaxies: fundamental parameters — galaxies: high-redshift — galaxies: structure

1. INTRODUCTION

The assembly history of nearby quiescent and morphologically early-type galaxies remains an important issue in extragalactic astronomy. Of particular interest is the fate of the population of compact red galaxies at redshift $z \sim 2$ (Daddi et al. 2005; Trujillo et al. 2006; van Dokkum et al. 2006) which has been the subject of much observational effort. To match the properties of local galaxies, the growth in size must be significantly larger than the growth in mass.

Although initial progress relied on photometric data, providing measures of both compact sizes and stellar masses of large samples beyond redshift $z \sim 1$ (e.g., Damjanov et al. 2009; Bezanson et al. 2011, 2012), key advances have become possible with spectroscopic samples. Spectroscopic data address the relative growth of the dynamical and stellar masses (Belli et al. 2014), as well as mean luminosity-weighted ages (Newman et al. 2014). As stellar velocity dispersions should remain stable through merger episodes, spectroscopic observations can link high-redshift progenitors with their local descendants. This is particularly important in considering progenitor bias, i.e., the continued arrival of recently-quenched larger galaxies (e.g., Carollo et al. 2013). Via the first comprehensive spectroscopic sample at $z > 1$, we quantified the size and mass growth rates of individual galaxies demonstrating significant growth (Belli et al. 2014) at a rate consistent with minor mergers observed from independent imaging studies (Newman et al. 2012).

Attention now focuses on understanding the population of massive compact sources at $z > 2$. Deep

imaging with the Wide Field Camera 3 (WFC3/IR) on-board *Hubble Space Telescope* (HST) has determined the growth rate is particularly rapid over the brief interval corresponding to $1.5 < z < 2.5$ (Newman et al. 2012). However, only limited spectroscopy is available as absorption line work is difficult in the near-infrared (e.g., Kriek et al. 2006). Until recently the relevant instruments (e.g. X-Shooter on the Very Large Telescope) were single-object long slit facilities. Despite heroic efforts, few stellar velocity dispersions are available beyond $z \simeq 2$ (van Dokkum et al. 2009; Toft et al. 2012; van de Sande et al. 2013).

The MOSFIRE multi-slit near-infrared spectrograph on Keck 1 (McLean et al. 2012) provides the first opportunity to systematically explore quiescent galaxies beyond $z \sim 2$. Here we present absorption line spectroscopy for a reasonable sample of compact massive galaxies at $2 < z < 2.5$. Our goal is to derive stellar velocity dispersions and dynamical masses, testing the rapid size growth rate inferred photometrically, as well as to examine this growth in the context of numerical simulations of galaxy merging.

Throughout we use AB magnitudes and assume a Λ CDM cosmology with $\Omega_M=0.3$, $\Omega_\Lambda=0.7$ and $H_0=70$ km s⁻¹ Mpc⁻¹.

2. DATA

2.1. Target Selection and Ancillary Data

To select spectroscopic targets, we use public photometric data from the NEWFIRM Medium-Band Survey (NMBS, Whitaker et al. 2011); this includes deep ground-based narrow, medium and broad band observations from the near-UV to the near-infrared and *Spitzer* IRAC and MIPS data. We limited our search to fields with high-quality HST F160W imaging

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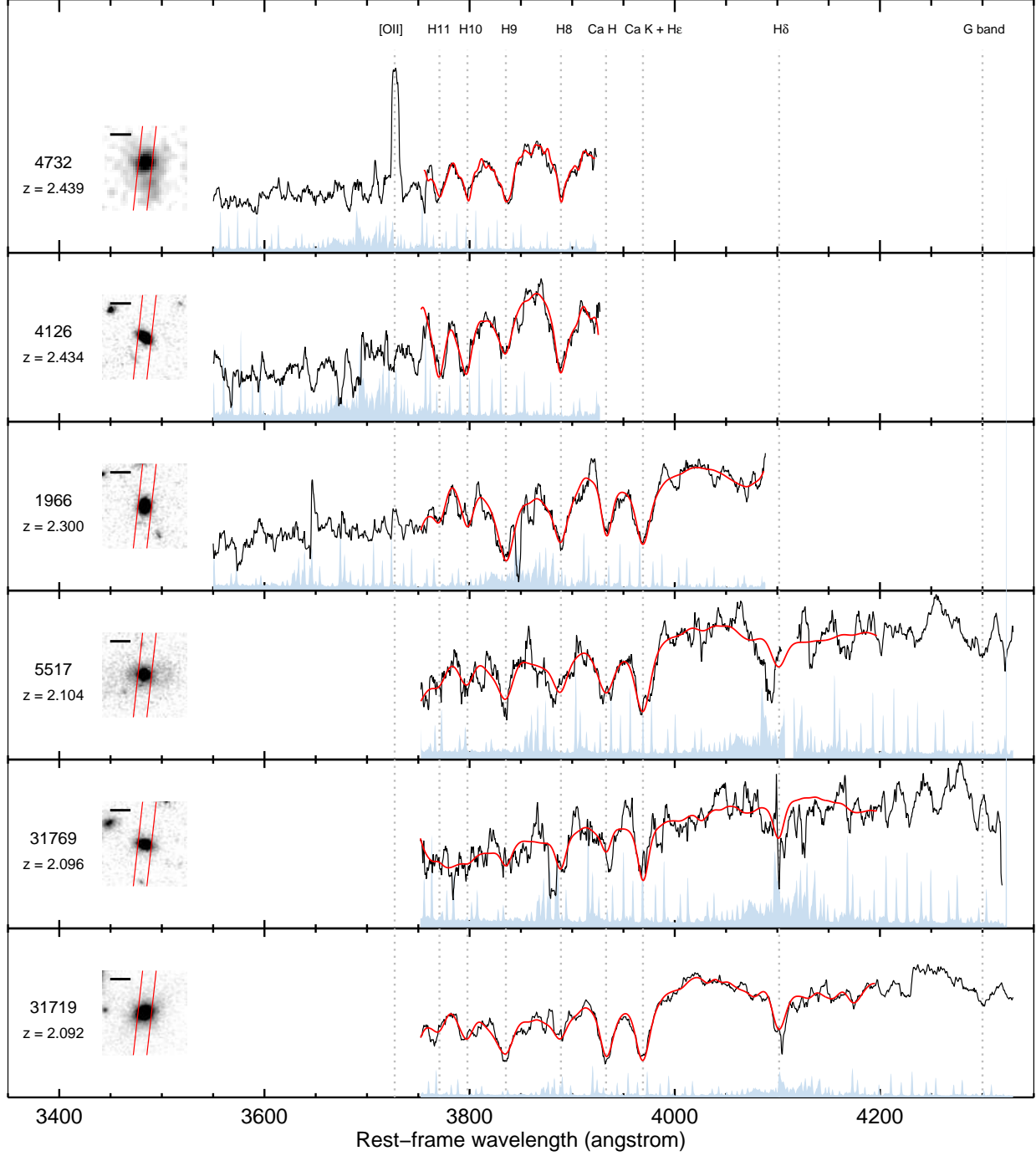


FIG. 1.— *HST* images and MOSFIRE spectra for a sample of six galaxies with detected absorption lines. For each object, the ID and spectroscopic redshift are indicated; the 4 arcsec cutout shows the F160W image with a 10 kpc ruler; the observed spectrum (inverse-variance smoothed, black line), its uncertainty (divided by 3, light blue), and the best-fit model (red line) are plotted. Absorption and emission features are marked by gray dotted lines. For 4732, only ground-based data are available, and the cutout is from the UltraVISTA H-band imaging (McCracken et al. 2012).

data from the Cosmic Assembly Near-IR Deep Extragalactic Legacy Survey (CANDELS, Grogin et al. 2011; Koekemoer et al. 2011).

We compiled a catalog of stellar masses, colors, and photometric redshifts derived via spectral energy distribution (SED) fitting. We selected the MOSFIRE pointing with the highest number of massive galaxies within the photometric redshift range $2 < z_{\text{phot}} < 2.5$, using $U - V$ and $V - J$ rest-frame colors to prioritize quiescent objects (Williams et al. 2009). Our best pointing in the COSMOS field includes 21 targets, some of which belong to the $z \sim 2.1$ protocluster discovered by Spitler et al. (2012).

2.2. Spectroscopic Data

Spectroscopic observations were undertaken using MOSFIRE on April 17, 2013, and March 6, 7, 2014 with clear sky and 0.5–0.7 arcsec seeing. We observed in the J band using a two-point dithering pattern and an exposure time of 120 s per frame, for a total of 8hr20min. The 0.8 arcsec slit width yielded a spectral resolution of 45 km/s. The data were reduced using the MOSFIRE pipeline that performs flat fielding, sky subtraction and wavelength calibration, and outputs rectified 2-D spectra, from which 1-D spectra were optimally extracted. Telluric correction and flux calibration were performed using the spectra of A0 standard stars.

Continuum emission is detected in many spectra and six, shown in Figure 1, exhibit absorption lines such as Ca II H and K and Balmer lines with well-defined redshifts in the range $2.092 < z < 2.439$ consistent with our photometric selection. Other features include the G band and the [OII] λ 3727, 3729 emission line. The continuum signal-to-noise ratios range from 5 to 18 per resolution element. Table 1 summarizes the observational data.

2.3. The Target Sample

To gain insight into the nature of our sample we use the UVJ diagram regularly employed to distinguish between star-forming and quiescent galaxies (e.g., Williams et al. 2009) (Figure 2). Rest-frame $U - V$ and $V - J$ colors were calculated by integrating the best-fit SEDs (see Section 3.2). We also show the $1 < z < 1.6$ sample of Belli et al. (2014) and the distinguishing line between quiescent and star-forming objects used in that work. Of the six selected galaxies, one (4732) is clearly star-forming given its [OII] emission, UVJ colors and strong $24 \mu\text{m}$ emission ($327 \pm 9 \mu\text{Jy}$). 4126 is likely recently-quenched: it lacks [OII] emission but shows strong Balmer absorption, much weaker $24 \mu\text{m}$ emission ($28 \pm 8 \mu\text{Jy}$), and UVJ colors of a post-starburst system. The other four galaxies show no signs of star formation: they have quiescent UVJ colors, no [OII] emission, and are not detected at $24 \mu\text{m}$ with the exception of 31769. For the latter, the mid-infrared emission ($161 \pm 11 \mu\text{Jy}$) is partly due to an active galactic nucleus (AGN), given its high X-ray luminosity ($2.0 \pm 0.3 \cdot 10^{44} \text{ erg s}^{-1}$, Elvis et al. 2009) and infrared colors (Stern et al. 2005). A more detailed study of the MOSFIRE spectra, including measurements of ages and star formation histories, will be presented in a future article (S. Belli et al., in prep.).

3. PHYSICAL PROPERTIES

3.1. Structural Properties

We derive structural parameters from the *HST* F160W data that probe the rest-frame optical emission. We fit 2D Sérsic profiles using GALFIT (Peng et al. 2002), masking out neighboring objects and deriving the point spread function (PSF) from isolated bright stars. For each object the fitting procedure gives the Sérsic index n , the axis ratio q , and the circularized effective radius $R_e = a\sqrt{q}$, where a is the half-light semi-major axis (see Table 1). From the tests performed by Newman et al. (2012), we estimate the uncertainty on the radii to be 10%.

The Sérsic profile fit is good for all objects except 5517, which presents an asymmetric halo perhaps because of a merging event. This object lies at the center of the protocluster “A” described by Spitler et al. (2012). We note that a luminous halo has been seen in at least another $z \sim 2$ protocluster member (Strazzullo et al. 2013). For object 5517 alone we estimate the effective radius by performing a curve of growth analysis on the radial profile, after deconvolving the image with the PSF. We tested this method on the other targets and find agreement with the Sérsic fit to within 30%.

The object 4732 lies just outside the area covered by CANDELS, and no size measurement was attempted.

3.2. Stellar Masses

We fit synthetic spectra to our extensive photometric data to measure stellar masses. We perform the fit using FAST (Kriek et al. 2009) adopting the Bruzual & Charlot (2003) template library and the Chabrier (2003) initial mass function (IMF). For the choice of the star formation history and stellar population parameters, we direct the reader to Belli et al. (2014). To ensure consistency between the mass and size measurements, for each object we scale the observed SED to match the F160W flux obtained via Sérsic fit, following Belli et al. (2014). The average correction is small, -0.06 ± 0.08 dex.

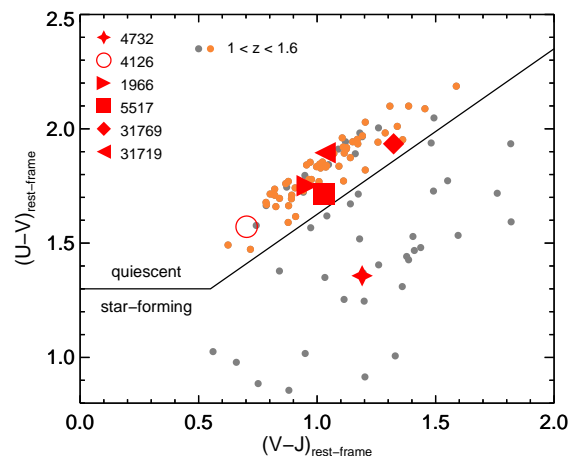


FIG. 2.— The UVJ diagram. Large symbols represent our MOSFIRE sample. Small circles are lower-redshift ($1 < z < 1.6$) galaxies of Belli et al. (2014), with orange indicating the quiescent subsample with velocity dispersion measurement (see Figures 3 and 4).

TABLE 1
PROPERTIES OF THE MOSFIRE SAMPLE

ID	R.A. (J2000)	Decl. (J2000)	J	$U - V$	$V - J$	z	σ_e (km s ⁻¹)	R_e (kpc)	n	q	$\log M_*/M_\odot$	$\log M_{\text{dyn}}/M_\odot$
31719	150.07393	2.2980	21.5	1.90	1.03	2.092	359 ± 30	3.08	5.5	0.79	11.62 ± 0.06	11.66 ± 0.09
31769	150.07460	2.3020	22.8	1.93	1.32	2.096	312 ± 65	1.98	5.4	0.54	11.28 ± 0.06	11.35 ± 0.19
5517	150.06562	2.2611	21.9	1.71	1.03	2.104	464 ± 138	5.07 ^a	11.41 ± 0.06	12.10 ± 0.17
1966	150.05489	2.1982	22.5	1.75	0.96	2.300	350 ± 61	1.01	2.7	0.71	11.22 ± 0.06	11.16 ± 0.16
4126	150.05579	2.2361	22.8	1.57	0.70	2.434	223 ± 56	1.25	1.4	0.45	10.89 ± 0.04	10.86 ± 0.22
4732	150.05246	2.2455	22.1	1.36	1.19	2.439	128 ± 48	11.40 ± 0.04	...

NOTE. — $U - V$ and $V - J$ are rest-frame colors. σ_e is the velocity dispersion within one effective radius. The effective radius R_e , Sérsic index n and axis ratio q are measured in F160W.

^a Size determined via curve of growth.

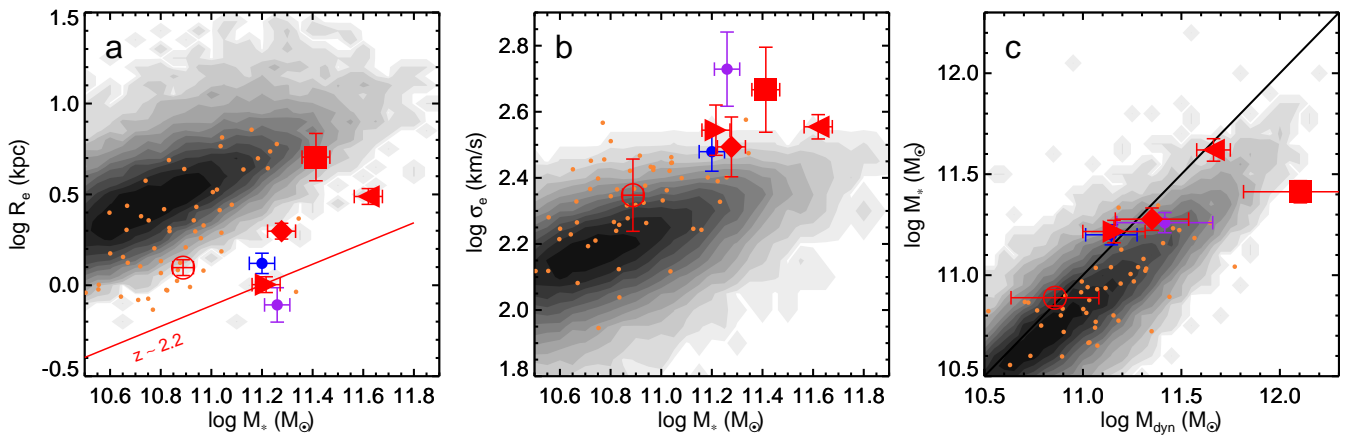


FIG. 3.— Relations between stellar mass and structural and dynamical properties for quiescent galaxies with $0 < z < 2.5$. Samples include: a local population from SDSS (grayscale), the $1 < z < 1.6$ sample from Belli et al. (2014, orange points), our MOSFIRE sample (red large symbols, as in Figure 2), and a compilation of $z > 2$ objects from the literature (van Dokkum et al. 2009, purple; van de Sande et al. 2013, blue). a) Mass-size relation. The red line indicates the $z \sim 2.2$ relation derived by Newman et al. (2012). b) Mass-velocity dispersion relation. c) Dynamical versus stellar mass. The black line indicates the one-to-one relation.

The corrected stellar masses M_* and their uncertainties are listed in Table 1. The masses, which range from 10^{11} to $10^{11.6} M_\odot$, are very large, as expected from the combination of selection criteria and bias due to the absorption line detection.

3.3. Velocity Dispersions

We measure the stellar velocity dispersions by fitting broadened templates to the MOSFIRE spectra in the range 3750–4200 Å. We use the Penalized Pixel-Fitting (pPXF) code of Cappellari & Emsellem (2004), and take template spectra from the library of synthetic stellar populations by Bruzual & Charlot (2003). Each spectrum is combined with an additive and a multiplicative polynomial to account for template mismatch and uncertainties in flux calibration or dust attenuation. Pixels contaminated by strong sky emission are excluded. The observed velocity dispersion is corrected for instrumental resolution, as measured from unblended sky lines, and template resolution.

We performed several tests to explore how robust are our derived velocity dispersions noting, in particular, the prevalence of Balmer absorption lines in some of the spectra. We followed the techniques discussed by van de Sande et al (2013) and Belli et al (2014) to which the

reader is referred. We considered variations in the utilized wavelength range, fitting polynomial degree, template models and fraction of discarded pixels. We also studied the template dependence by comparing our results with those based on the higher resolution Indo-US library (Valdes et al 2004). For each object, the systematic error was determined by adding the contribution of these effects in quadrature to the random error. Finally, to achieve the dispersions within the effective radius σ_e reported in Table 1, we apply a 5% aperture correction (based on the model by van de Sande et al. 2013).

Object 31719 has already been observed with X-Shooter by van de Sande et al. (2013), that derived $\sigma_e = 446^{+54}_{-59}$ km s⁻¹. Our measurement is smaller, with a discrepancy at the 1.4σ level. In the following, we will use our measurement for this object, since our spectrum has a higher signal-to-noise ratio.

4. THE DYNAMICAL AND STRUCTURAL EVOLUTION OF QUIESCENT GALAXIES

We use our new dynamical measurements, together with those from lower-redshift observations, to constrain the evolution of the size and structure of quiescent galaxies.

Figure 3a shows the mass-size relation for similarly-

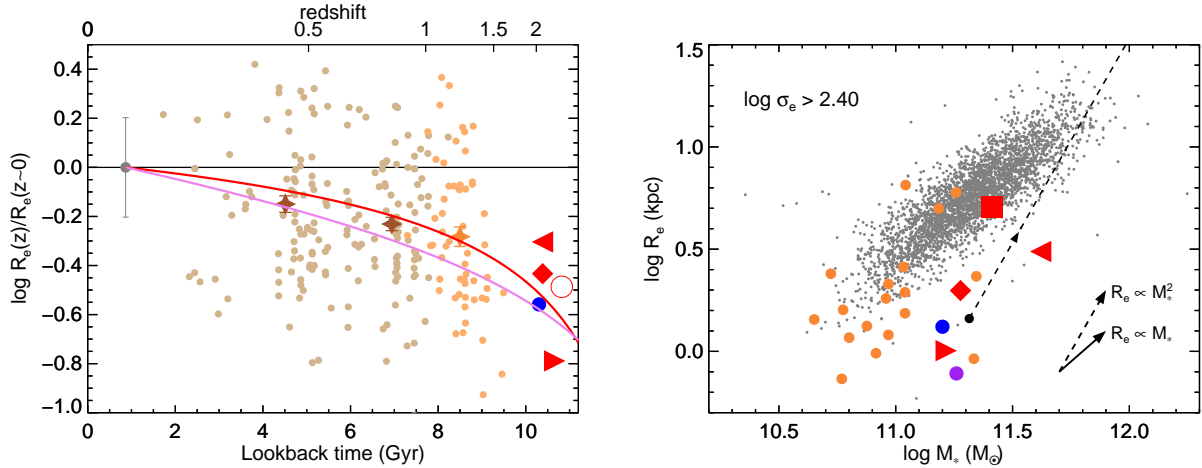


FIG. 4.— Left: redshift evolution of effective radius at fixed velocity dispersion. The small points are a compilation of lower-redshift studies (Treu et al. 2005, brown, Belli et al. 2014, orange), with the diamonds representing the running average. The red line is the fit to the evolution of the mass-size relation from Newman et al. (2012), while the violet line is a fit to the observed evolution at fixed number density from van Dokkum et al. (2010). Right: Mass-size plane for $0 < z < 2$ quiescent galaxies with $\log \sigma_e > 2.40$. The high-redshift sample is clearly offset from the local population (gray points) toward small radii, at fixed stellar mass. Symbols as in Figures 2 and 3. The black point represents the mean of the high-redshift sample, and an evolutionary track with slope $\alpha = 2$ is shown as a dashed line. The solid and dashed arrows represent the cases for $\alpha = 1$ and 2 respectively.

selected UVJ -quiescent galaxies over $0 < z < 2.5$. We show the local population from the Sloan Digital Sky Survey (SDSS DR7, Abazajian et al. 2009, grayscale map), a sample at $1 < z < 1.6$ from Belli et al. (2014, small orange points), and the other two quiescent galaxies at $z > 2$ for which velocity dispersion measurements have been published (van Dokkum et al. 2009, purple, and van de Sande et al. 2013, blue; see also Toft et al. 2012). Our MOSFIRE sample is shown in red. Clearly, the mass-size relation evolves with redshift. At fixed stellar mass, galaxies at $z \sim 1.3$ are about 0.25 dex smaller than the local population. At $z > 2$, the logarithmic offset from the local sample nearly doubles, implying evolution accelerates at earlier cosmic times. These findings confirm the results of previous photometric studies, as shown by the agreement of our points with the mass-size relation at $z \sim 2.2$ from Newman et al. (2012, red line). Object 5517, the brightest galaxy (BCG) of a protocluster, is an exception, suggesting that such systems have large sizes already at $z \sim 2$, in agreement with other studies at $z \lesssim 1.8$ (Papovich et al. 2012; Stanford et al. 2012; Newman et al. 2014).

The main advance of this paper is that we can now explore the dynamical properties of galaxies at $z > 2$. Figure 3b shows high-redshift galaxies have significantly larger velocity dispersions than lower redshift objects of similar stellar mass. Most of our MOSFIRE objects have similar dispersions, $\sigma_e \approx 300 \text{ km s}^{-1}$, with the exception of 4126 (shown as an open symbol). This galaxy is the only post-starburst object in our sample, and presents an elongated morphology and low Sérsic index, $n = 1.4$, typical of disk-like galaxies.

Velocity dispersions enable us to calculate dynamical masses, via $M_{\text{dyn}} = 5\sigma_e^2 R_e / G$. Figure 3c compares the dynamical and stellar masses, again contrasting the trend with samples drawn from the literature. While the $z \sim 1.3$ sample closely follows the local distribution, $z > 2$ galaxies tend to have higher stellar-to-dynamical mass ratios. This difference was first suggested by Toft et al. (2012) and van de Sande et al. (2013), but as our MOS-

FIRE sample doubles the number of dynamical masses with $z > 2$, it is now more significant, particularly as our velocity dispersions are more accurate.

We calculate the average mass ratio for all the $z > 2$ quiescent galaxies, excluding the BCG, and we find $\log M_*/M_{\text{dyn}} = -0.02 \pm 0.03$ and an average stellar mass of $10^{11.3} M_\odot$. Although our sample is modest, this is a remarkably tight agreement. Considering the most massive galaxies of the $1 < z < 1.6$ sample, we find $\log M_*/M_{\text{dyn}} = -0.23 \pm 0.05$ and an average mass of $10^{11.1} M_\odot$. This significant evolution could arise if $z > 2$ quiescent galaxies have a reduced dark matter fraction, a heavier stellar IMF or different structure compared to their lower-redshift counterparts.

We now use our dynamical masses to infer the rate of size growth for high-redshift quiescent galaxies. Following Belli et al. (2014), we assume that we can link progenitor and descendant galaxies by selecting populations at *fixed velocity dispersion*. This follows the results of numerical simulations that show that velocity dispersion is minimally affected during merger events (e.g. Hopkins et al. 2009; Oser et al. 2012), and the observed unchanging velocity dispersion function (Bezanson et al. 2012). For each of our $z > 2$ objects, we therefore select SDSS galaxies with similar velocity dispersions (within 0.05 dex). We calculate the median size of the local subsample and assume that our high redshift object will grow in size until it reaches that value. We repeat this procedure for lower-redshifts samples and show the inferred growth rates in the left panel of Figure 4. Typical $z > 2$ galaxies³ are noticeably smaller than those at $z \sim 1.3$, despite the small interval in cosmic time between the two samples. Our results are qualitatively in agreement with the size growth at *fixed number density* derived by van Dokkum et al. (2010, violet line). Interest-

³ The BCG and the van Dokkum et al. (2009) galaxy are omitted because in the local universe there are no objects with such high velocity dispersions. Additionally both have the most uncertain dispersions.

ingly, the growth of individual galaxies is not dissimilar to that of the total population derived from the evolution of the mass-size relation (Newman et al. 2012, red line). Although minor merging is likely responsible for the size evolution of quiescent galaxies at $z < 1.5$, such a rapid growth at $z \sim 2$ is hard to reconcile with the observed merger rate (Newman et al. 2012; see also Cimatti et al. 2012).

In the right panel of Figure 4 we show the mass-size relation for galaxies with $\log \sigma_e > 2.40$ at all redshifts. This velocity dispersion bin includes all the $z > 2$ galaxies except the post-starburst object 4126. The local population forms a clear sequence whereas objects at intermediate redshifts show a mild offset towards smaller sizes (and masses). However, the $z > 2$ galaxies populate a distinct region of the mass-size plane with almost no overlap with lower redshift samples; almost all have radii 0.3 dex below the local sample. The inescapable conclusion is that quiescent galaxies at $z > 2$ must undergo a dramatic and rapid size growth.

A powerful method to constrain the physical processes responsible for this size growth is to measure the slope $\alpha = d \log R_e / d \log M_*$ of the evolutionary tracks on the mass-size plane and compare it with theoretical predictions. Simple virial arguments (Naab et al. 2009; Bezanson et al. 2009) give $\alpha = 1$ for identical mergers and $\alpha = 2$ for the limiting case of mergers with infinitely diffuse satellites. More realistic numerical simulations, which include the effect of dark matter, gas, and a distribution of orbits, indicate that minor mergers are less efficient than the theoretical limit, and yield values in the range $1.4 < \alpha < 1.8$ (Hopkins et al. 2009; Nipoti et al. 2012; Oser et al. 2012; Posti et al. 2014). The simulations of Hilz et al. (2013), in which massive dark matter halos enhance the efficiency of minor merging up to $\alpha = 2.4$, are the only exception. However, the large dark matter fraction at the center of these simulated galaxies disagrees with the observed stellar-to-dynamical mass ra-

tios at both low and high redshift.

Assuming evolution at fixed velocity dispersion, we measure α by considering the tracks that high-redshift points must follow in order to match the local distribution. Using this technique, the $z \sim 1.3$ sample yields $\alpha = 1.4 \pm 0.2$ (Belli et al. 2014). Merging can therefore readily explain the size growth over $0 < z < 1.5$, a conclusion supported by direct imaging (Newman et al. 2012). However, at $z > 2$ the growth is clearly much more rapid. It is not possible to derive a robust measurement of the slope α for two reasons: firstly, our $z > 2$ sample is not velocity dispersion complete; secondly, one of our basic assumptions might not hold, since the velocity dispersion function has not been probed beyond $z \sim 1.5$. As the number density of quiescent galaxies declines steeply at this redshift, a strong progenitor bias is expected. Despite these limitations, we can still derive an important lower limit on α , by assuming that the average high-redshift galaxy (shown in black, excluding the BCG) will evolve into one of the most massive objects at $z \sim 0$. Using this method, for all the $z > 2$ quiescent galaxies excluding the BCG we derive a lower limit of $\alpha \gtrsim 2$, shown as a dashed line in the figure.

In summary, our spectroscopic data allows us to conclude that both the absolute rate of size growth *and* the inferred motion in the mass-size plane are independently inconsistent with minor mergers being the principal physical process governing the evolution of quiescent galaxies at $z \sim 2$.

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REFERENCES

- Abazajian, K. N., et al. 2009, *ApJS*, 182, 543
 Belli, S., Newman, A. B., & Ellis, R. S. 2014, *ApJ*, 783, 117
 Bezanson, R., van Dokkum, P., & Franx, M. 2012, *ApJ*, 760, 62
 Bezanson, R., van Dokkum, P. G., Tal, T., Marchesini, D., Kriek, M., Franx, M., & Coppi, P. 2009, *ApJ*, 697, 1290
 Bezanson, R., et al. 2011, *ApJ*, 737, L31
 Bruzual, G., & Charlot, S. 2003, *MNRAS*, 344, 1000
 Cappellari, M., & Emsellem, E. 2004, *PASP*, 116, 138
 Carollo, C. M., et al. 2013, *ApJ*, 773, 112
 Chabrier, G. 2003, *PASP*, 115, 763
 Cimatti, A., Nipoti, C., & Cassata, P. 2012, *MNRAS*, 422, L62
 Daddi, E., et al. 2005, *ApJ*, 626, 680
 Damjanov, I., et al. 2009, *ApJ*, 695, 101
 Elvis, M., et al. 2009, *ApJS*, 184, 158
 Grogin, N. A., et al. 2011, *ApJS*, 197, 35
 Hilz, M., Naab, T., & Ostriker, J. P. 2013, *MNRAS*, 429, 2924
 Hopkins, P. F., Hernquist, L., Cox, T. J., Keres, D., & Wuyts, S. 2009, *ApJ*, 691, 1424
 Koekemoer, A. M., et al. 2011, *ApJS*, 197, 36
 Kriek, M., van Dokkum, P. G., Labbé, I., Franx, M., Illingworth, G. D., Marchesini, D., & Quadri, R. F. 2009, *ApJ*, 700, 221
 Kriek, M., et al. 2006, *ApJ*, 649, L71
 McCracken, H. J., et al. 2012, *A&A*, 544, A156
 McLean, I. S., et al. 2012, in *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 8446, Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series
 Naab, T., Johansson, P. H., & Ostriker, J. P. 2009, *ApJ*, 699, L178
 Newman, A. B., Ellis, R. S., Andreon, S., Treu, T., Raichoor, A., & Trinchieri, G. 2014, *ApJ*, in press
 Newman, A. B., Ellis, R. S., Bundy, K., & Treu, T. 2012, *ApJ*, 746, 162
 Nipoti, C., Treu, T., Leauthaud, A., Bundy, K., Newman, A. B., & Auger, M. W. 2012, *MNRAS*, 422, 1714
 Oser, L., Naab, T., Ostriker, J. P., & Johansson, P. H. 2012, *ApJ*, 744, 63
 Papovich, C., et al. 2012, *ApJ*, 750, 93
 Peng, C. Y., Ho, L. C., Impey, C. D., & Rix, H.-W. 2002, *AJ*, 124, 266
 Posti, L., Nipoti, C., Stiavelli, M., & Ciotti, L. 2014, *MNRAS*, 440, 610
 Spitler, L. R., et al. 2012, *ApJ*, 748, L21
 Stanford, S. A., et al. 2012, *ApJ*, 753, 164
 Stern, D., et al. 2005, *ApJ*, 631, 163
 Strazzullo, V., et al. 2013, *ApJ*, 772, 118
 Toft, S., Gallazzi, A., Zirm, A., Wold, M., Zibetti, S., Grillo, C., & Man, A. 2012, *ApJ*, 754, 3
 Treu, T., et al. 2005, *ApJ*, 633, 174
 Trujillo, I., et al. 2006, *ApJ*, 650, 18
 Valdes, F., Gupta, R., Rose, J. A., Singh, H. P., & Bell, D. J. 2004, *ApJS*, 152, 251
 van de Sande, J., et al. 2013, *ApJ*, 771, 85
 van Dokkum, P. G., Kriek, M., & Franx, M. 2009, *Nature*, 460, 717

van Dokkum, P. G., et al. 2006, ApJ, 638, L59
—, 2010, ApJ, 709, 1018
Whitaker, K. E., et al. 2011, ApJ, 735, 86

Williams, R. J., Quadri, R. F., Franx, M., van Dokkum, P., &
Labbé, I. 2009, ApJ, 691, 1879